

OH spectral evolution of oxygen-rich late-type stars

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Received date; accepted 27 January 2004

Abstract. We investigated the main-line spectral evolution with shell thickness of oxygen rich AGB stars. The study is based on a sample of 30 sources distributed along the IRAS colour-colour diagram. The sources were chosen to trace the Miras with thick shells and the whole range of OH/IR stars. The Miras exhibit a 1665 MHz emission strength comparable to that at 1667 MHz. Even though the Miras of the study have quite thick shells, their spectral characteristics in both main lines attest to a strong heterogeneity in their OH shell with, in particular, the presence of significant turbulence and acceleration. The expansion velocity has been found to be about the same at 1665 and 1667 MHz, taking into account a possible velocity turbulence of 1-2 km s⁻¹ at the location of the main-line maser emission. An increase in the intensity ratio 1667/1665 with shell thickness has been found. A plausible explanation for such a phenomenon is that competitive gain in favour of the 1667 MHz line increases when the shell is getting thicker. There is an evolution in the spectral profile shape with the appearance of a substantial inter-peak signal when the shell is getting thicker. Also, inter-peak components are found and can be as strong as the external standard peaks when the shell is very thick. This trend for an increase of the signal in between the two main peaks is thought to be the result of an increase of the saturation with shell thickness. All sources but two - a Mira and an OH/IR star from the lower part of the colour-colour diagram - are weakly polarized. The strong polarization observed for those two particular objects is thought to be the result of perturbations in their shells.

Key words. Stars: AGB and post-AGB – circumstellar matter – Masers – Line: profiles – Stars: mass-loss – Stars: evolution

1. Introduction

The standard model proposed by Reid et al. (1977) explains convincingly the double-peaked profile usually observed for Miras and OH/IR objects, particularly at 1612 MHz. It also explains the general features of the circumstellar envelopes. Nevertheless, many stars do not have such a simple OH maser profile. In particular, in the main lines, the peaks are often broad and consist of many components. To explain these more complex profiles new models have been suggested.

Alcock & Ross (1986) have shown that it is possible to change the profile shape for a partially or fully saturated maser when the masering region becomes sufficiently thick. When that happens, the profiles consist of broader peaks and a signal is seen at the velocity of the star. This is in agreement with the spectral profiles observed for Miras and OH/IR objects, but the inferred internal and external radii disagree with the MERLIN and VLA interferometric observations. To overcome this discrepancy, Alcock & Ross

suggest that the mass loss is not a continuous and uniform process but rather consists of ejection of blobs of material in random directions such that the symmetrical geometry is preserved. The model calculation results of Collison & Nedoluha (1995) agree with the conclusions by Alcock & Ross that the observed OH profile cannot be produced by the standard model description of a smooth, spherically symmetric, steady windlike mass loss. The assumption of a non-smooth shell is justified by the interferometric maps themselves. Indeed, with the increase in interferometer resolution the more detailed maps revealed that the OH maser emission is located in clumps (Welty, Fix & Mutel, 1987; Chapman, Cohen & Saikia, 1991) and/or the shell shows evidence for deviation from spherical symmetry (Diamond et al. 1985; Bowers, Johnston & de Vegt 1989).

Evidence for acceleration in the OH shell has been found for many stars (Chapman et al. 1994; Etoka & Le Squeren 1996, 1997; Szymczak et al. 1998; Richards et al. 1999). The importance of acceleration in the shape of the spectral profile is underlined by the work of Chapman & Cohen (1985). According to the value of the logarithmic velocity gradient ϵ , the shape of the observed line profile

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largely varies as follows : (1) for $\epsilon \rightarrow 0$ the classical double-peaked profile is observed; (2) for $0 < \epsilon < 1$ an inter-peak signal can be seen; (3) for $\epsilon \simeq 1$ a plateau-shaped profile is obtained and finally (4) for $\epsilon > 1$ the profile is reduced to a single peak centred on the stellar velocity. Thus, many spectral profiles resembling those observed can be modelled by introducing acceleration in the OH shell.

Based on the IRAS measurements at 12, 25 and 60 μm , Olnon et al. (1984) discovered the existence of a continuous sequence from Miras to OH/IR objects in the so-called colour-colour diagram which plots the colour indices $[60-25]$ vs. $[12-25]$. This sequence has been analysed in terms of shell thickness by van der Veen & Habing (1988) and van der Veen & Rutgers (1989). These studies lead to a commonly accepted interpretation that it is an evolutionary track followed by intermediate-mass stars. In this scenario, Miras are the progenitors of OH/IR stars involving among other characteristics an increase of their mass-loss rate and expansion velocity with time. Nonetheless, the initial mass of the star could play a role in the setting up of that sequence. Based on an alternative model for the evolution of AGB stars proposed by Epchtein et al. (1990), Lépine et al. (1995) proposed that the IRAS colour-colour sequence could be a sequence of mass.

The study presented here deals with the spectral profiles in both main lines of Miras and OH/IR stars distributed along the colour-colour diagram. The main aim is to determine whether there is a relation between the spectral profile and the thickness of the circumstellar shell. The next section presents the observations and selection criteria. Section 3 and Appendices A, B and C present the results concerning the individual stars and a comparison with previous observations. The discussion is presented in Sect. 4 and the final conclusions in Sect. 5.

2. Observations

The OH observations were carried out between March and December 1994 with the Nançay transit radiotelescope. The half-power beamwidth of the radiotelescope at 1.6 GHz is $3.5'$ in α by $19'$ in δ . The system temperature was about 50 K. A frequency switching mode was used. The autocorrelator was divided into 4 banks of 256 channels each, allowing simultaneous observations of the two OH main lines (i.e., 1665 and 1667 MHz) in both left- and right-handed circular polarizations (LHC, RHC). A bandwidth of 1.56 MHz was used, leading to a velocity resolution of 0.56 km s^{-1} . The ratio of flux to antenna temperature was 1.1 Jy K^{-1} at 0° declination. Each observation lasted about one hour providing a mean rms of 0.035 K. All the radial velocities given hereafter are relative to the Local Standard of Rest (LSR).

2.1. Selection criteria

All the sources presented here had already been positively detected at 1667 MHz and a few of them at 1665 MHz too. We looked for 1665 MHz emission for a possible signature

of circumstellar shell evolution. Figure 1 shows the colour-colour diagram of the selected sources. The colour indices $[25-12]$ and $[60-25]$ used here were calculated according to the following relation :

$$[\lambda_1 - \lambda_2] = \log\left(\frac{\nu_1 S_\nu(\lambda_1)}{\nu_2 S_\nu(\lambda_2)}\right) \quad (1)$$

similar to that used by Sivagnanam et al. (1990).

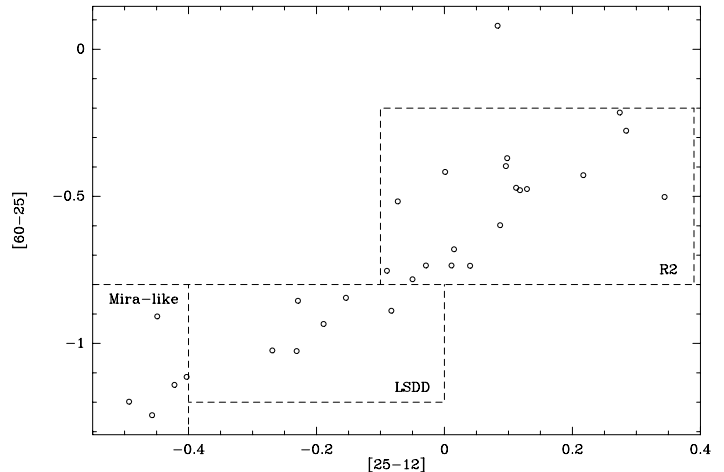


Fig. 1. $[60-25]$ vs. $[25-12]$ colour-colour diagram of the sources of the study

The selected objects were intended to cover the far right-hand side of the region labelled “Mira-like” by Sivagnanam et al. (1990) and the areas labelled LSDD and R_2 by David et al. (1993a) for the OH/IR sources. These areas are respectively delimited as follows :

$$\begin{aligned} -0.50 < [25-12] < -0.40 \\ -1.30 < [60-25] < -0.80 \end{aligned} \quad (2)$$

$$\begin{aligned} -0.40 < [25-12] < +0.00 \\ -1.20 < [60-25] < -0.80 \end{aligned} \quad (3)$$

and

$$\begin{aligned} -0.10 < [25-12] < +0.40 \\ -0.80 < [60-25] < -0.20 \end{aligned} \quad (4)$$

These areas are covered by the regions labelled IIIa (i.e., *variable stars with more evolved O-rich circumstellar shells*), IIIb (i.e., *variable stars with thick O-rich circumstellar shells*) and IV (i.e., *variable stars with very thick O-rich circumstellar shells*) by van der Veen & Habing (1988) and van der Veen & Rutgers (1989).

The observations by David et al. (1993a) were done at 1612 and 1667 MHz, and those of Sivagnanam et al. (1990) at 1667 MHz only. Both sets of observations were

carried out with the Nançay radiotelescope with the same resolution as the present work, allowing a direct comparison of the profile shapes. Observations made by David et al. (1993a) for the stars of interest were taken over a period of 5.5 years from mid-1986 to the early 1992. Those of Sivagnanam et al. (1990) were all obtained in July 1986 except that of IRAS 20547+0247 which was obtained in May 1987.

3. Results: General comments

Table 1 presents the main characteristics of the sources of the study. Columns 1 to 10 give the name of the source, its galactic coordinates, its LRS class, its IRAS flux at 12, 25 and 60 μm , its [25 – 12] and [60 – 25] colour indices and finally the name under which it is also known. The table was divided such that the three parts correspond to the three regions “Mira-like”, LSDD and R_2 defined by the colour indices given in (2), (3) and (4) respectively.

The results concerning the spectral profile changes of the individual sources in the three regions of concern are presented in Appendix A.

Figure B.1 presents the spectral profiles of each source at 1665 and 1667 MHz in both circular polarizations. As IRAS 18198–1249, IRAS 18257–1000, IRAS 18488–0107, IRAS 19065+0832, IRAS 22177+5936 and IRAS 19075+0921 had already been detected before the work of David et al. (1993a), their spectra were not displayed by the previously cited authors. The spectral profiles of the 5 first cited sources are given in Dickinson & Turner (1991) and that of IRAS 19075+0921 is given in Le Squeren et al. (1992). Therefore, for these 6 sources, the spectral shape changes were determined from the comparison of our observations with those of the two papers mentioned, while for the other sources, comparison has been made with those of David et al. for the OH/IR stars and Sivagnanam et al. (1990) for the Miras. All the sources of the R_2 region were also observed at 1612 MHz by David et al. (1993b) and quite a few by te Lintel et al. (1991). Three, out of the six sources of the LSDD region were observed at 1612 MHz by te Lintel et al. (1991). Finally, all the sources of the “Mira-like” region were observed at 1667 and 1612 MHz by Sivagnanam et al. (1990).

Tables C.1 and C.2 present the results of the spectral measurements done at 1665 and 1667 MHz respectively. As for Table 1, these tables were divided into three parts corresponding to the three regions of concern. Columns 1 to 16 give the name of the source, a comment about its profile shape in the case of a non “standard” double-peaked profile, the velocity of the blue- and red-shifted peaks and the velocity of any inter-peak component in both circular polarizations, the stellar and expansion velocities inferred, the flux intensity of the blue- and red-shifted peak and that of any inter-peak component in both circular polarizations (i.e., measured respectively

at the velocity given in column 3 to 8) respectively. For the profile shape we distinguish between 4 different categories : “standard” type spectrum but exhibiting inter-peak signal (“s”), plateau profile spectrum (“p”), spectrum with more than two well detached peaks (i.e., “standard” profile plus an internal feature, “ci” where “internal” means in between the two standard peaks) and finally spectrum showing only a single peak centred at or about the stellar velocity (“c”). For the two last categories, the characteristics of the signal are given in columns 7, 8 and 15, 16.

David et al. (1993a) used the 1985 IRAS fluxes while we used the corrected 1987 fluxes. Thus, IRAS 18460-0254 which was initially in the R_2 region of David et al. is now outside this box due to a notable difference in its [60 – 25] colour index when calculated with or without correction (cf. Fig. 1). Nevertheless, its [25 – 12] colour index is hardly affected by the application of the IRAS correction. Since this index is the most relevant in terms of mass loss and shell thickness dependence, we decided to incorporate this source in the statistics of the R_2 region.

4. Discussion

4.1. The “Mira-like” sample

The area corresponding to the OH-emitter Miras contains five objects which, according to their location in the colour-colour diagram, have the thickest shells of their category. None of the five sources exhibits a standard spectral profile either at 1665 or 1667 MHz and only one source shows a double-peaked profile characterized by two unevenly broad and spiky peaks. This source has the smallest [25–12] colour index value of the sample. All the sources show quite irregular spectral profile shapes. This tendency is also clearly observable in the spectral profiles obtained by Sivagnanam et al. (1990, their Fig. 2). Such a strong departure from the double-peaked spectral profile indicates a strong heterogeneity in the physical and dynamical conditions ruling over the whole OH shell even though those Miras have quite thick shells.

All the sources show OH variability. Among the five sources selected, four of them were primarily successfully detected in the satellite line at 1612 MHz and we detected 1665 MHz emission from all of them. For most of the sources of the “Mira-like” sample positively detected by Sivagnanam et al. (1990) as OH-emitters, when 1612 MHz emission was detected its intensity was comparable to or greater than that at 1667 MHz. Such a trend in the strengthening of the 1612 MHz emission over that of the main lines is expected when the shell is getting thick enough (cf. Lewis 1989; David *et al.* 1993b; Habing 1996). The sample selected here (i.e., $-0.50 < [25-12] < -0.40$) is such that it fulfils the physical conditions for the 1612 MHz and the main-line emissions to be of similar strength. IRAS 15226–3603 was one of the rare sources not detected positively at 1667 MHz, nor did we detect any

Table 1. Galactic positions, IRAS flux at 12, 25 and 60 μm for the sources of the study

Irass name	l deg.	b deg.	LRS class	$S_{12\mu\text{m}}$ Jy	$S_{25\mu\text{m}}$ Jy	$S_{60\mu\text{m}}$ Jy	[25-12]	[60-25]	other name
15226+3603	334.7	17.0	28	166.40	121.10	16.58	-0.457	-1.244	
15262+0400	8.1	45.8	27	46.15	30.87	4.70	-0.493	-1.198	
20396+0826	38.3	-28.4	69	11.71	9.23	1.60	-0.422	-1.141	XX Aqr
20547+0247	51.4	-26.1	32	45.52	33.76	10.02	-0.449	-0.908	U Equ
22525+6033	109.2	1.1	24	112.10	92.30	17.05	-0.403	-1.114	
17482+2824	1.1	-0.8	42	428.70	480.40	109.10	-0.269	-1.024	OH 1.1-0.8
17505+3143	358.5	-2.9	34	47.60	69.51	23.82	-0.154	-0.845	
19075+0921	43.4	0.2	12	133.00	163.70	54.91	-0.229	-0.855	
19161+2343	57.1	5.1	31	112.20	137.20	31.04	-0.231	-1.026	
19190+1128	46.6	-1.3	32	27.13	36.60	10.22	-0.189	-0.934	
19244+1115	47.1	-2.5	28	1346.00	2314.00	717.70	-0.083	-0.889	IRC 10420
17220+2448	1.0	6.1		5.35	9.06	3.84	-0.090	-0.753	
17317+3331	354.9	-0.5	79	104.00	291.50	234.50	0.129	-0.475	
17385+3332	355.6	-1.7		2.88	13.25	10.01	0.344	-0.502	
17392+3319	355.9	-1.7	37	19.12	37.29	16.49	-0.029	-0.735	
17411+3154	357.3	-1.3		1262.00	2723.00	1365.00	0.015	-0.680	
17418+2713	1.4	1.0	79	14.95	51.28	45.94	0.217	-0.428	OH 1.3+1.0
17550+2120	8.0	1.4		5.38	21.08	30.81	0.274	-0.215	OH 8.0+1.4
18033+2111	9.1	-0.2		9.75	20.84	9.20	0.011	-0.735	
18135+1456	15.7	0.8	79	31.02	124.40	157.60	0.284	-0.277	OH 15.7+0.8
18198+1249	18.3	0.4		15.56	42.54	33.91	0.118	-0.479	OH 18.3+0.4
18257+1000	21.5	0.5	39	46.32	120.50	115.90	0.096	-0.397	OH 21.3+0.4
18348+0526	26.5	0.6		359.80	633.80	463.00	-0.073	-0.517	OH 26.5+0.6
18432+0149	30.7	0.4		25.06	52.32	48.09	0.001	-0.417	OH 30.7+0.4
18460+0254	30.1	-0.7		111.10	279.90	806.90	0.083	0.080	OH 30.2-0.7
18488+0107	31.9	-0.4		16.45	42.98	44.03	0.098	-0.370	OH 32.0-0.5
19065+0832	42.6	0.1	39	20.22	51.52	31.18	0.087	-0.598	OH 42.6+0.0
19254+1631	51.8	-0.2	39	16.57	44.67	36.27	0.112	-0.471	OH 51.8-0.2
19352+2030	56.4	-0.3	03	42.89	97.88	43.12	0.040	-0.736	OH 56.4-0.3
22177+5936	104.9	2.4	38	123.20	228.90	90.65	-0.050	-0.782	OH 104.9+2.4

1665 MHz emission from it. Since this source was detected positively by Sivagnanam *et al.* earlier, our non-detection is most likely due to the fact that the source was at its OH minimum at the time of our observations. Finally, one source (IRAS 22525+6033) appeared to be strongly polarized in both main lines. This source exhibits the strongest 1665 MHz emission and the greatest expansion velocity of its group: about 10 km s^{-1} . Thus, this source is located in the narrow zone where both Miras and OH/IR stars can be found in the “period–OH expansion velocity” diagram as given by Sivagnanam *et al.* (1989).

4.2. The OH/IR sample of the LSDD region

This sample of six sources presents the highest rate of double-peaked spectral profiles at 1667 MHz (>50%) with very faint or non-existent inter-peak emission. Four of the six sources (i.e., 67%) were detected at 1665 MHz at the time of the observations. Nevertheless, none of the sources shows a “classical” double-peaked profile at that frequency. In addition, the 1665 MHz emission for the sources of this sample is very faint. Only IRAS 19075+0921 shows a significant flux density at 1665 MHz (i.e., comparable to that observed at 1667 MHz). This particular source also shows strong polarization in both main lines and has the smallest expansion velocity of the whole sample, with a value of 7.7 km s^{-1} . Consequently, like IRAS 22525+6033, IRAS 19075+0921 is located in the zone where Miras and OH/IR stars coexist in the “period–OH expansion velocity” diagram of Sivagnanam *et al.* (1989).

4.3. The OH/IR sample of the R_2 region

The R_2 sample is composed of nineteen OH/IR stars. 84% of the stars of this group show at least two peaks in their 1667 MHz spectral profile with only 31.5% showing a standard double-peaked profile (i.e., with no inter-peak signal). The percentage of stars having at least two peaks at 1665 MHz is only 48%. Five sources (i.e., 26%) exhibit a single-peaked spectral profile and five other sources did not show any 1665 MHz emission at all. 26% of the sources that were positively detected at 1665 MHz show an internal group of components with an intensity comparable or superior to that of the red- or blue-shifted peaks.

Among the whole set of the double-peaked profiles of this group of thicker OH/IR stars, nine sources (i.e., 48%) show an inter-peak signal at 1667 MHz. Sun (1991) has modelled the evolution of the 1612 MHz double-peaked profile when the mass-loss rate varies with time. The author shows that the inter-peak signal at 1612 MHz increases with the mass-loss rate. Our results show that this property also applies to the 1667 MHz emission. Also, the existence of an inter-peak signal is related to the degree of saturation (Fix 1978). Therefore, the increased number of inter-peak signals observed in the profiles of these thicker OH/IR stars is quite probably a signature of greater maser saturation at 1667 MHz.

4.4. Spectral profile shape as a function of the shell thickness

Considering the three groups together, it stands out that the spectral profiles with inter-peak signal as strong as that of the red- or blue-shifted peaks are only observed

for OH/IR sources at 1665 MHz. OH/IR stars with inter-peak signal at 1667 MHz are found in both the LSDD and the R_2 regions. Except IRAS 19352+2030 which shows very faint blue- and red-shifted peaks and a very strong inter-peak component in both main lines, the sources show a faint 1667 MHz inter-peak signal which can be slightly polarized. Spectral profiles showing a well-detached inter-peak component at 1665 MHz arise from OH/IR stars belonging to the upper part of the R_2 region, that is, from the stars which have a very thick shell. The “plateau” profiles are only observed at 1665 MHz, but in all three groups. It can be particularly pronounced for some objects (cf. IRAS 17411–3154). Such a profile shape generally attests to significant acceleration (i.e., Chapman & Cohen, 1985 for a velocity gradient $\epsilon \simeq 1$). This confirms that the 1665 MHz emissive zone indeed originates deeper inside the OH shell than the 1667 MHz zone.

Figure 2 shows the correlation between the expansion velocity inferred at 1665 MHz and that at 1667 MHz. This diagram uses the measurements from all the sources that showed at least two peaks in both main lines except IRAS 19244+1115 for which the expansion velocities inferred have to be taken with caution (cf. Appendix A.2). The total number of such sources is 12, that is about 41% of the whole sample of 29 objects for this statistic. This corresponds respectively to 40, 40 and 42% of the “Mira-like”, LSDD and R_2 samples. The dotted line in the figure is the locus for $V_{1665 \text{ MHz}} = V_{1667 \text{ MHz}}$. Most of the sources have a similar expansion velocity in both main lines. At the location of the OH shell, some 10^{16} cm away from the central star, turbulence of $1\text{--}2 \text{ km s}^{-1}$ is expected (Deguchi, 1982). This could result in the broadening of the line wings and therefore explain most of the cases where 1665 MHz profiles overshoot those of the 1667 MHz line by such a value. One source, IRAS 17392–3319 from the R_2 region, shows an expansion velocity at 1665 MHz greater by 4 km s^{-1} than that inferred at 1667 MHz. This may be due to the faintness of the signal at 1665 MHz leading to a bigger uncertainty in the exact location of the peak maxima.

Only four sources show an intensity at 1665 MHz greater than that at 1667 MHz. These sources belong to the “Mira-like” and LSDD regions. None of the sources of the R_2 region show this peculiarity. We have investigated whether there was any correlation between the intensity observed in the main lines and the thickness of the shell. Taking all the sources showing both 1665 and 1667 MHz emission, we have found that there is a relation between the ratio of the intensity observed in the two main lines and the $[25\text{--}12]$ colour index. This relation is shown Fig. 3, where the dotted line shows the linear regression over all the sources except the three for which $\text{Intensity}_{1667} \gg$

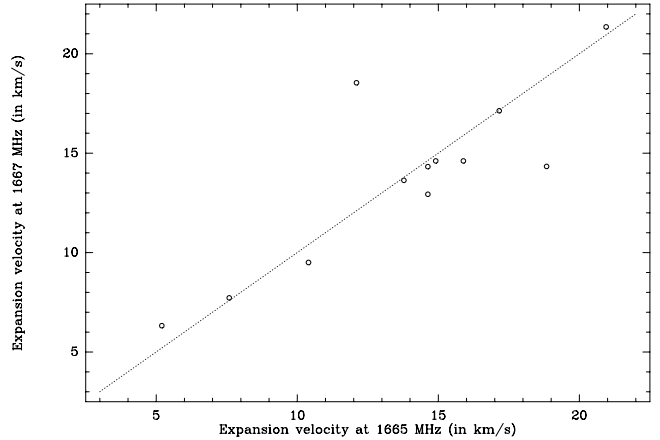


Fig. 2. Expansion velocity at 1667 MHz versus that at 1665 MHz

Intensity_{1665} (plotted with filled symbols). The linear regression is given by :

$$\text{Int}(1667/1665)_i = 6.54 \times [25 - 12]_i + 4.40 \quad (5)$$

where $\text{Int}(1667/1665)_i$ is the ratio of the maximum intensity measured at 1667 and 1665 MHz for the source “i” and $[25\text{--}12]_i$ is its colour index.

Three sources do not lie on the relation given by Eq. (5). We do not think that this could be due to variability. Indeed, for one of the sources (IRAS 18432–0149) the 1665 MHz emission faintness ($\sim 2\sigma$) could be responsible for the high $\text{Int}(1667/1665)$ found but such an explanation is not plausible for the two other sources which showed reasonably strong 1665 MHz emission at the time of the observations. On the other hand, it is noteworthy that the three sources not following the general trend have similar $[25\text{--}12]$ indices : $[0; 0.1]$.

Equation (5) shows that when the shell becomes thicker, the 1667 MHz emission is favoured in comparison with that at 1665 MHz. Since the expansion velocity is about the same at 1665 and 1667 MHz it is expected that the maser emission from both main lines should be found at similar distances from the star and possibly share the same regions of the OH shell. This hypothesis is supported by the maps of late-type stars obtained in VLBI (cf. Chapman et al. 1991, 1994; Szymczak et al. 1999). Therefore, a plausible explanation for such a trend is an increase of competitive gain in favour of the 1667 MHz line when the shell becomes thicker.

This increase of the ratio of the 1667 MHz emission over that at 1665 MHz is in agreement with the fact that 1667 MHz emission becomes relatively stronger and eventually dominant in the post-AGB stage (Zijlstra et al. 1989, 2001). The dominance of the 1667 MHz emission at this stage of the stellar evolution is supported by theoretical studies (Field 1985, Field & Gray 1988).

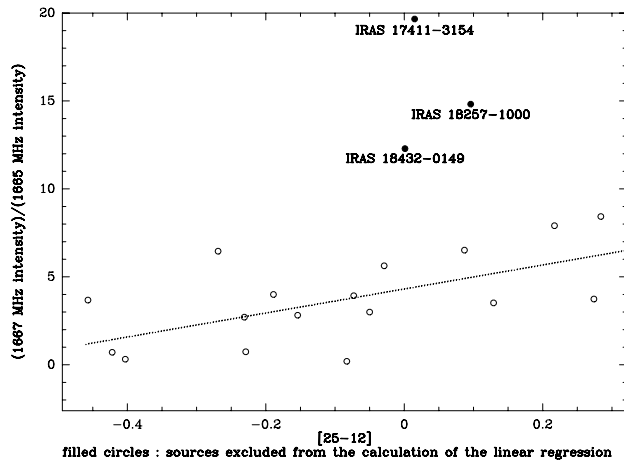


Fig. 3. (1667 MHz)/(1665 MHz) intensity ratio versus the colour index [25-12]

4.5. The Mira-OH/IR sequence

The Miras-OH/IR sequence seen in the IRAS colour-colour diagram results from the combination of various relations existing amongst the physical parameters of the central star and the circumstellar envelopes : (1) between the mass-loss rate and the period (Vassiliadis & Wood 1993), (2) between the period and the expansion velocity (Dickinson & Chaison 1973, Le Squeren et al. 1979), (3) between the OH luminosity and the flux at 25 μ m (Sivagnanam et al. 1989), (4) between the OH radius and the OH luminosity (Baud & Habing 1983), (5) between the size of the maser regions and the mass-loss rate (predicted for the H₂O masers by Cooke & Elitzur 1985 and observed by Cohen 1987, Lane et al. 1987 and Yates & Cohen 1994) and (6) between the mass-loss rate and the optical depth at 10 μ m (Bedijn 1988). This can be summarized as follows : the increase of the pulsation period, the expansion velocity, the radius of the circumstellar shell and the luminosity are all function of the mass-loss rate and the opacity of the shell.

The sequence observed in the colour-colour diagram may be either evolutionary or related to the stellar mass. If it is an evolutionary sequence the star should experience an increase of its mass-loss rate with time. If it is a mass sequence the mass loss should be connected somehow to the initial mass. Volk & Kwok (1988) studied the IR spectral evolution of AGB stars. Their study shows that AGB stars that have a silicate signature at 10 μ m in absorption are mainly stars with an initial mass on the main sequence greater than 3 M_{\odot} . This implies that the shell thickness depends on the initial mass, since only the stars massive enough will have a thick shell i.e. will become OH/IR stars. Also, through a study of the galactic distribution of OH/IR stars according to their expansion velocity, Likkell (1989) found that the reddest OH/IR objects tend to have the highest initial masses.

This latter observational result is in agreement with earlier work by Baud (1981) who showed the existence of a relation between the expansion velocity and the initial mass of the central star.

We found that IRAS 22525+6033 and IRAS 19075-0921 belonging to the “Mira-like” group and to the LSDD group respectively show very similar characteristics : 1665 MHz flux comparable to that at 1667 MHz, strong polarization over the whole velocity range of the spectrum in both main lines, similar OH expansion velocity values and [25-12] indices. This puts these two objects in the location where Miras and OH/IR stars coexist both in the “period-OH expansion velocity” diagram and the colour-colour diagram, but they have a completely unusual set of spectral characteristics in comparison with the rest of the stars with similar expansion velocities and IRAS [12-25] fluxes. In particular, the very strong polarization observed over the whole spectral profile is clearly a sign of perturbation undergone throughout the whole OH circumstellar shell, which we interpret to be a signature that these objects are in the transition phase Mira \rightarrow OH/IR star. Indeed, similar polarization characteristics in the OH shell have been observed (Bains et al. 2003, Etoka & Diamond 2004) attesting to the importance of the role played by the magnetic field (in particular in the envelope shaping) in the transitional stages of evolved stars. Bearing in mind the existing relations between the initial mass and some of the actual physical properties of AGB stars such as their shell thickness or their expansion velocity, it is very likely that the IRAS colour-colour sequence is related to both the initial mass of the object and its evolutionary stage. A likely hypothesis is that the object arrives more or less high on the sequence and evolves through a part of it before expelling its shell entirely via an increasing mass loss.

5. Conclusion

Acceleration and turbulence can strongly alter the spectral profile shapes of circumstellar OH masers. This is well illustrated by the variety of spectral profiles observed for the “Mira-like” sample. When the shell gets thicker and less turbulent, the standard double-peaked profile settles. At that stage, saturation is the main cause of spectral profile alteration at 1667 MHz, causing an inter-peak signal to appear while plateau-shape or strong inter-peak components observed in the spectral profiles at 1665 MHz still attest to the presence of acceleration. Even though the spectral profile characteristics show that the 1665 MHz emission can be found deeper in the OH shell than the 1667 MHz emission, the expansion velocity is found to be about the same in both main lines, taking into account a possible velocity turbulence of 1-2 km s⁻¹ at the location of the main-line maser emission. It is then expected that the maser emission from both main lines should be found at similar distances from the star and possibly share

the same regions of the OH shell. Also, the intensity ratio 1667/1665 has been found to increase with the shell thickness. This can be explained as an increase of competitive gain in favour of the 1667 MHz when the shell is becoming thicker. Finally, a Mira and an OH/IR star from the lower part of the colour–colour diagram show quite strong polarization while the rest of the sources are weakly polarized. Strong polarization can be the sign of dramatic changes in the shell itself. We suspect that the strong polarization observed in those two particular objects may be linked with their being in the phase of transition Mira \rightarrow OH/IR star.

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Appendix A: Remarks on individual sources

A.1. “Mira-like”

IRAS 15226–3603 : We did not detect any emission from this source either at 1665 or at 1667 MHz in March 1994 although Sivagnanam et al. (1990) clearly detected a signal at 1667 MHz in July 1986 with an intensity of 0.46 Jy for the strongest spectral profile component. At the time of the observations by Sivagnanam et al. this source was not detected at 1612 MHz, thus this Mira is certainly a type I.

IRAS 15262+0400 : At the time of our observations this source showed an intensity at 1667 MHz seven times greater than that observed by Sivagnanam et al. (1990) in July 1986. Nevertheless, the intensity ratio $I_{\text{red peak}}/I_{\text{blue peak}}$ (where “red peak” and “blue peak” stand for the red-shifted and blue-shifted emission respectively) stayed the same. We detected faint non-polarized emission ($\simeq 0.2$ K) at 1665 MHz internal in velocity to that observed at 1667 MHz (where “internal (in velocity)” should be understood as defined in Sect. 3, meaning in this case that the signal is located between the blue and red peaks observed at 1667 MHz).

IRAS 20396–0826 : This identified Mira, also known as XX Aqr, has an optical period of 334 days (Kholopov et al. 1985, *GCVS*). It was classified by Sivagnanam et al. (1990) as a “peculiar object” due to the difference between the spectral profile shape observed at 1612 and that observed at 1667 MHz. The 1612 MHz emission shows a double-peaked profile covering a velocity range of 6 km s^{-1} . The 1667 MHz emission observed by Sivagnanam et al. is internal in velocity to that observed at 1612 MHz and has a plateau-shaped profile. Our observations show that the 1665 MHz emission is internal in velocity to that observed at 1667 MHz with a single peak centred on the stellar velocity. A strong acceleration ($\epsilon > 1$) at the location of the 1665 MHz emission could produce such a spectral profile shape (cf. Chapman & Cohen 1985), strengthening the hypothesis of a velocity gradient in the OH shell of this source as proposed by Sivagnanam et al.

IRAS 20547+0247 : This source, also known as U Equ, was indexed by Sivagnanam et al. as the Mira having the smallest expansion velocity with a value of 2.4 km s^{-1} (cf. their Table 1). The 1667 MHz intensity of this source has substantially dropped. The blue-shifted peak which was already very faint at the time of the observations by Sivagnanam et al. in May 1987 is no longer detected in our observations. We detected a very faint 1665 MHz signal only in the red-shifted peak centred at the same velocity as the emission we observed at 1667 MHz.

IRAS 22525+6033 : Sivagnanam et al. classified this source as a “peculiar object” due to its spectral

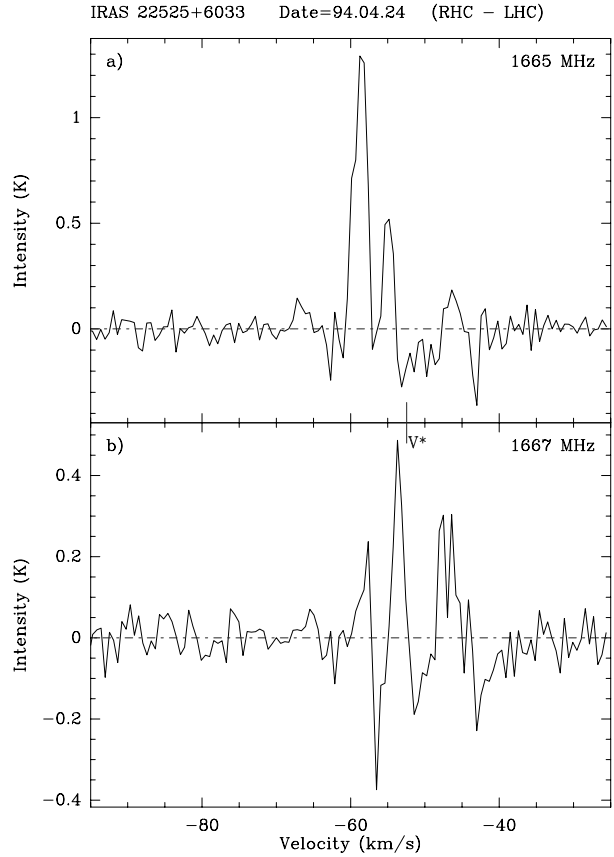


Fig. A.1. Stokes V spectra of IRAS 22525+6033: **a)** at 1665 MHz, **b)** at 1667 MHz. V^* is the stellar velocity inferred from the present work

profile at 1667 MHz. Our observations indeed confirm the non-standard profile shape in both main lines, with multiple components quite well detached most of which are strongly right-hand polarized (cf. Fig A.1 showing the Stokes V spectra for both main lines). The resulting spectral profiles in the two circular polarizations are therefore very different (cf. Fig B.1). Sivagnanam et al. inferred an expansion velocity of 4.4 km s^{-1} . From our observations we infer an expansion velocity more than twice that value : $9\text{--}10 \text{ km s}^{-1}$.

A.2. OH/IR stars from the LDSS region

IRAS 17482–2824 : Emission at 1667 and 1665 MHz is present. The velocity location for the red- and blue-shifted peaks is about the same in both lines but the 1667-MHz red-shifted peak emission overshoots that at 1665 MHz by about 1 km s^{-1} . Even though there is a separation of two years between the 1667-MHz observations made by David et al. (1993a) and ours, there is no significant change in the spectral profile and intensity ratio $I_{\text{red peak}}/$

$I_{\text{blue peak}}$.

IRAS 17505–3143 : The comparison between the 1667-MHz observations made by David et al. (1993a) and ours showed a noticeable change in both the flux and the intensity ratio $I_{\text{red peak}}/I_{\text{blue peak}}$ between the two epochs of observations. We did not detect any 1665 MHz emission. This source was observed repeatedly at 1612 MHz. In particular, te Lintel et al. (1991) observed it in November 1985. They found an expansion velocity of 15 km s^{-1} with the blue- and red-shifted peaks located at -12.2 and $+19.2 \text{ km s}^{-1}$ respectively (cf. their Table 2). From our 1667 MHz observation we found the blue- and red-shifted peaks to be centred at -12.9 and $+20.3 \text{ km s}^{-1}$ respectively and an expansion velocity of 16.6 km s^{-1} in agreement with David et al. (1993a). This implies that the 1667 MHz red-shifted peak overshoots that at 1612 MHz by about 1 km s^{-1} .

IRAS 19075+0921 : The spectral profile of this source shows detached multi-components in both main lines leading to a non-standard shape, particularly noticeable at 1665 MHz. Like IRAS 22525+6033, this source shows a strong degree of polarization. Nevertheless, in this object the left-handed polarization prevails as can be seen in Fig. A.2 which displays the Stokes V spectra in both main lines. As a result of the strong degree of circular polarization of some of the spectral components, the profile is quite different in both circular polarizations for both main lines. The 1667 MHz spectrum profile obtained by Le Squeren et al. (1992) in January 1986 is similar to ours. The intensity ratio $I_{\text{red peak}}/I_{\text{blue peak}}$ did not show any significant change between the two sets of observations. We also found the same characteristic velocities as those given by Le Squeren et al. (1992, their Table 1).

IRAS 19161+2343 : We detected only a very faint and narrow emission in the red-shifted peak at 1665 MHz. The 1667 MHz red-shifted peak overshoots that at 1665 MHz by $\sim 1 \text{ km s}^{-1}$. No significant change in the profile shape can be seen between the observations made by David et al. (1993a) in May 1990 and ours.

IRAS 19190+1128 : This source shows very faint and asymmetric emission at 1667 MHz. In particular, the blue-shifted peak hardly reached a 2σ detection. We do not notice any significant change in the profile shape between the observations made by David et al. (1993a) in August 1986 and ours, performed 8 years later. We did not detect any emission at 1665 MHz.

IRAS 19244+1115 : This source, also known as IRC 10420, has quite a broad spectral profile. Because of the protocol of our observations the red-shifted part of the spectra has been cut. Indeed, if we refer to the observations made by David et al. (1993a) at 1667 MHz, we can clearly see the existence of a strong red-shifted com-

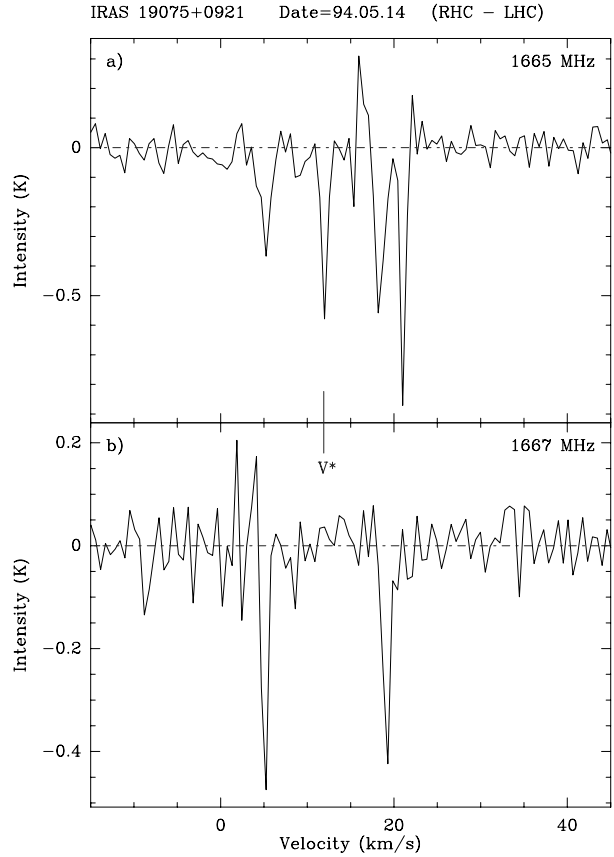


Fig. A.2. Stokes V spectra of IRAS 19075+0921: **a)** at 1665 MHz, **b)** at 1667 MHz

ponent centred at $+110 \text{ km s}^{-1}$, outside our bandwidth coverage. Although it was covered by our bandwidth, we did not detect the most blue-shifted component centred at $+41 \text{ km s}^{-1}$ seen in the profile published by David et al.. That is most probably due to the “band edge effects” concerning the first and last 5 to 10 channels of the band. Due to the inadequate protocol of observation for this particular source, the stellar and expansion velocities given in Tables C.1 and C.2 should be treated with caution. This source was not used in the statistical study on the expansion velocity presented in Sect. 4.4.

A.3. OH/IR stars from the R_2 region

IRAS 17220–2448 : We did not detect any emission at 1667 MHz, but a faint, narrow left-hand polarized component centred at $V=+45.3 \text{ km s}^{-1}$ was detected at 1665 MHz. Considering the velocity characteristics given by David et al. (1993a, $V_*=+39.4 \text{ km s}^{-1}$ and $V_{exp}=15.4 \text{ km s}^{-1}$) this component should be an OH maser signature of an inner region belonging to the back part of the shell.

IRAS 17317–3331 : The 1667 MHz emission of this source exhibits an inter-peak signal. At 1665 MHz we observed an internal group of components between the two asymmetrical red- and blue-shifted peaks. This internal peak is centred about the stellar velocity and its intensity is of the same order as that observed for the blue-shifted peak. This source was also observed at 1612 MHz by te Lintel et al. (1991) and by David et al. (1993a). From our observations in the main lines, we found the same stellar velocity as David et al. which is in agreement with that given by te Lintel et al. considering their spectral resolution. At 1612 MHz the spectral profile does not have any noticeable particularity. Both te Lintel et al. and David et al. classified it in the “standard double-peaked profile” category.

IRAS 17385–3332 : Both main lines exhibit faint emission. The blue-shifted peak is not detected at 1667 MHz and only just detected at 1665 MHz. At 1665 MHz, we clearly observe an internal peak centred on the stellar velocity, and the bluest narrow component we detected is centred at the same velocity as the 1612 MHz blue-shifted peak as given by te Lintel et al. (1991) and David et al. (1993b).

IRAS 17392–3319 : We did not detect any emission at 1665 MHz. The 1667 MHz emission shows a standard double-peaked profile. The profile and intensity have not shown any noticeable change between our observations carried out in 1994 and those performed by David et al. 8 years earlier.

IRAS 17411–3154 : The 1665 MHz emission exhibits a non-standard profile shape where at least 8 well-defined components are observed, of which 4 are coming from strong internal emission. Interestingly, the back part of the shell (i.e., $V > V_*$) shows quite strong right-handed polarization while the front part (i.e., $V < V_*$) is left-hand polarized (Fig. A.3). A similar phenomenon has already been observed in other sources (i.e., Ukita & Le Squeren 1984; Szymczak et al. 2001), and can be explained by the Cook mechanism when the magnetic field is aligned on the line of sight (Cook 1966, 1977). The 1667 MHz spectral profile is more typical although showing inter-peak emission. This source also exhibits strong 1612 MHz emission with intensity reaching a value as high as 200-300 Jy (cf. Table 2 of te Lintel et al. 1991 and Table 1 of David et al. 1993b). The expansion velocity inferred at 1612 MHz is of the same order as that inferred at 1667 MHz (i.e., 18.5 km s⁻¹, Table C.2). On the other hand, the most red-shifted part of the 1665 MHz emission was not detected, leading to a much smaller expansion velocity in that line.

IRAS 17418–2713 : This source shows a single component at 1665 MHz centred about $V = -21$ km s⁻¹, that is about the mid-point between the stellar velocity and the 1667 MHz blue peak velocity. The expansion velocity

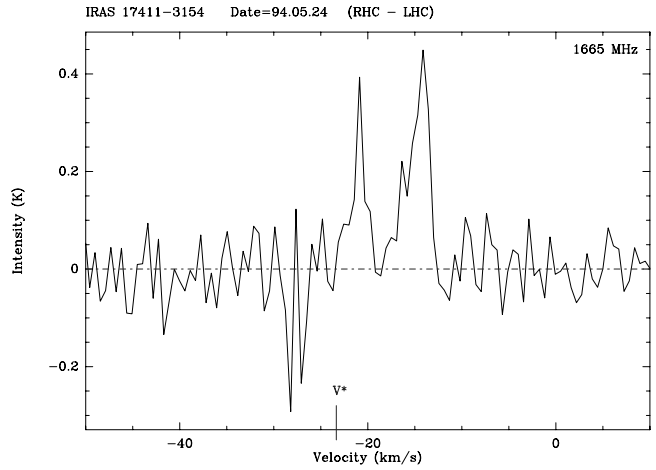


Fig. A.3. Stokes V spectrum of IRAS 17411–3154 at 1665 MHz.

inferred at 1667 MHz is more than 1 km s⁻¹ greater than that inferred by David et al. (1993b) and te Lintel et al. (1991) from the 1612 MHz emission. This is due to an overshoot of the 1667 MHz blue peak over that at 1612 MHz.

IRAS 17550–2120 : Emission is observed in both main lines but with an inverse intensity ratio $I_{\text{red peak}}/I_{\text{blue peak}}$. The 1667 MHz profile underwent a strong change of shape since 1987 when a triple-peaked profile shape was observed (cf. Fig. 2 of David et al. 1993a) while we detected only the two standard peaks. The extra internal component centred about $V = -17$ km s⁻¹ detected by David et al. vanished between the two epochs of observation. The nature of this event is unclear since the source has not been monitored.

IRAS 18033–2111 : It seems that the OH emission reported by David et al. (1993a, their Fig. 2 and Table 1) corresponds to two distinct sources : one emitting in the velocity range $[-130;-100]$ km s⁻¹ (the one we observed) and the second emitting in the velocity range $[-70;-30]$ km s⁻¹. The 1612 MHz observations reported by te Lintel et al. (1991) and David et al. (1993b) deal with the velocity range $[-70;-30]$ km s⁻¹. To our knowledge no observation has been done at 1612 MHz in the velocity range $[-130;-100]$ km s⁻¹. At the date of our observation neither 1667 MHz nor 1665 MHz emissions were detected while David et al. (1993a) clearly reported emission at 1667 MHz. Thus, the source corresponding to the velocity range $[-130;-100]$ km s⁻¹ is clearly variable.

IRAS 18135–1456 : The 1667 MHz spectrum of this source shows an asymmetric double-peaked profile. The 1665 MHz spectrum is internal to that at 1667 MHz and emission is observed over the whole velocity range (leading to its classification as a plateau-shaped profile).

Actually, the two external components (i.e., the bluest and reddest Doppler shifted) are those with the faintest emission. In particular, if one excepts the strong component centred at $V=-9.4 \text{ km s}^{-1}$, the general shape of the profile resembles a gaussian. This could be the signature of an acceleration zone (cf. Chapman & Cohen 1985).

IRAS 18198–1249, IRAS 18257–1000, IRAS 18488–0107 and IRAS 19065+0832 : These four sources were observed by Dickinson & Turner (1991) in both main lines. All four sources show a standard double-peaked profile at 1667 MHz. We did not notice any significant change in their spectral profiles. We did not detect any 1665 MHz either for IRAS 18198–1249 or for IRAS 18488–0107. We positively detected IRAS 18257–1000 and IRAS 19065+0832 at 1665 MHz. Dickinson & Turner classified IRAS 18198–1249, IRAS 18257–1000 and IRAS 18488–0107 as sources for which the 1667 MHz emission overshoots that at 1612 MHz.

IRAS 18348–0526 : This source shows a strong inter-peak emission at 1667 MHz. At 1665 MHz we also noted a non-zero inter-peak emission whose shape resembles a very flat-topped gaussian centred on the stellar velocity. This kind of shape is similar to that presented by Bowers (1991) for an isotropic outflow in a spherical circumstellar shell with significant rotation. This hypothesis would imply a non-negligible rotation component in the external regions of the circumstellar shell considering the radius expected for the OH formation (Goldreich & Scoville 1976, Huggins & Glassgold 1982). This is very unlikely. Another more credible hypothesis regards the emission as the sum of two components. The first component, coming from an external OH region expanding radially with an acceleration such as $1 > \epsilon > 0$, would give the standard red- and blue-shifted peaks with possibly a faint inter-peak signal. The second component would come from an inner OH strongly accelerating zone with $\epsilon > 1$, given the internal gaussian-like emission observed (cf. Chapman & Cohen 1985, their Fig. 3). This scenario is sustained by the 1667 MHz profile shape which also shows an inter-peak signal, strengthening the hypothesis that some acceleration is still present at the location of the main-line maser emission.

IRAS 18432–0149, IRAS 18460–0254 and IRAS 19254+1631 : The first two sources show faint emission at 1665 MHz barely above the noise threshold. The 1667 MHz spectra for all three sources did not show any significant change between the observations carried out by David et al. (1993a) and ours.

IRAS 19352+2030 : At 1667 MHz we detected emission from three peaks centred at $V=-8.70$, -1.4 and $+7.0 \text{ km s}^{-1}$. The two most red- and blue-shifted peaks are very faint. Only the central peak at $V=-1.40 \text{ km s}^{-1}$ was detected by David et al. (1993a). We also detected

emission at 1665 MHz centred at $V=-1.40 \text{ km s}^{-1}$. This peak is significantly left-hand polarized (i.e., 15-20%) in both main lines. This source was also observed at 1612 MHz by David et al. (1993b). They only detected a single peak centred at $V=-3.22 \text{ km s}^{-1}$.

IRAS 22177+5936 : At 1667 MHz this source shows a double-peaked profile with strong emission between the the red- and blue-shifted peaks. At 1665 MHz only the blue-shifted peak is clearly detected. The red-shifted peak is barely above the noise threshold in right-handed polarization. The spectral profiles obtained by Dickinson & Turner (1991) are similar to ours for both main lines. This source was also observed at 1612 MHz by te Lintel et al. (1991). The characteristic velocities they found are in agreement with ours.

Appendix B: Spectral line atlas

Appendix C: Spectral measurement results

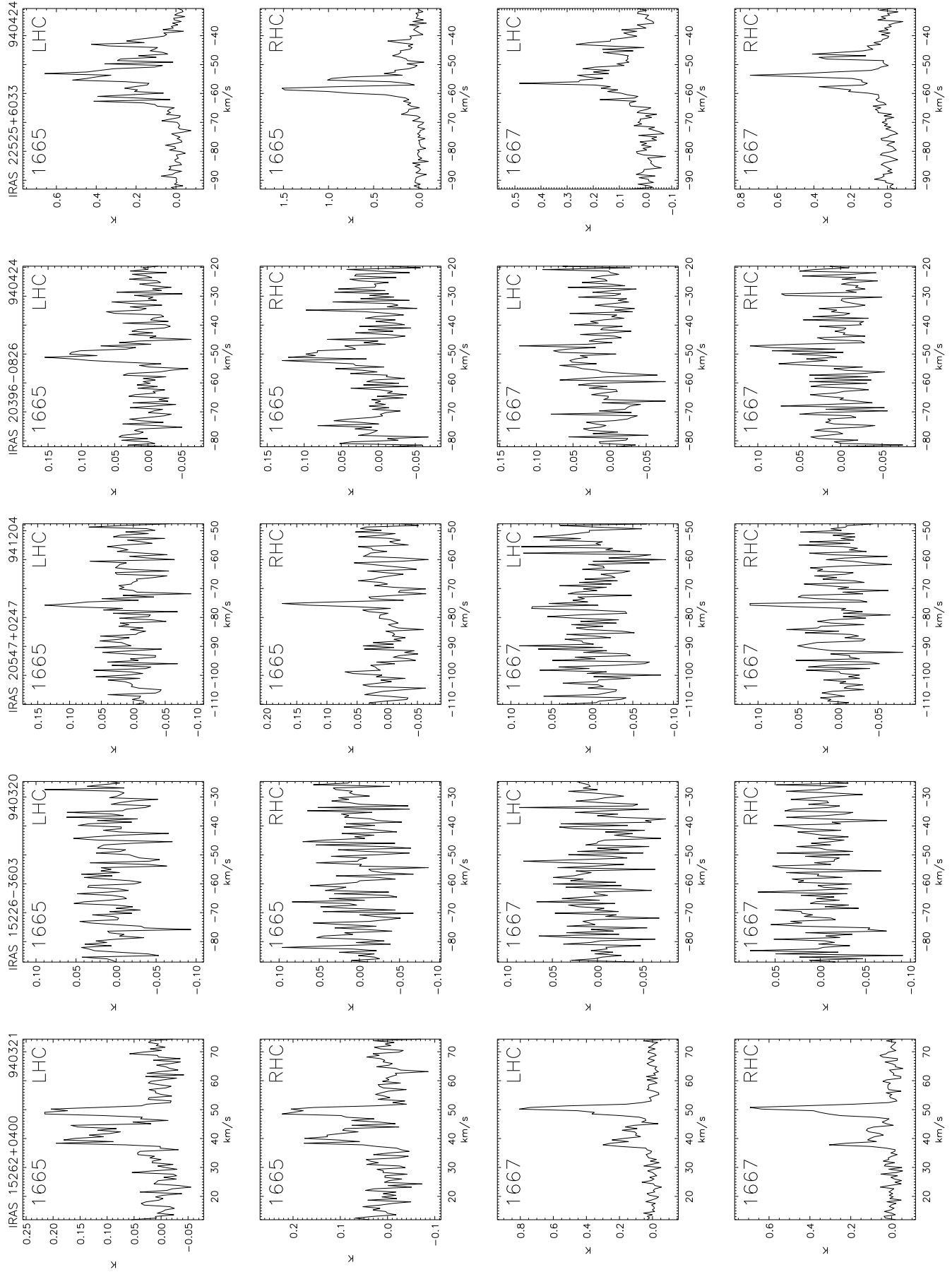
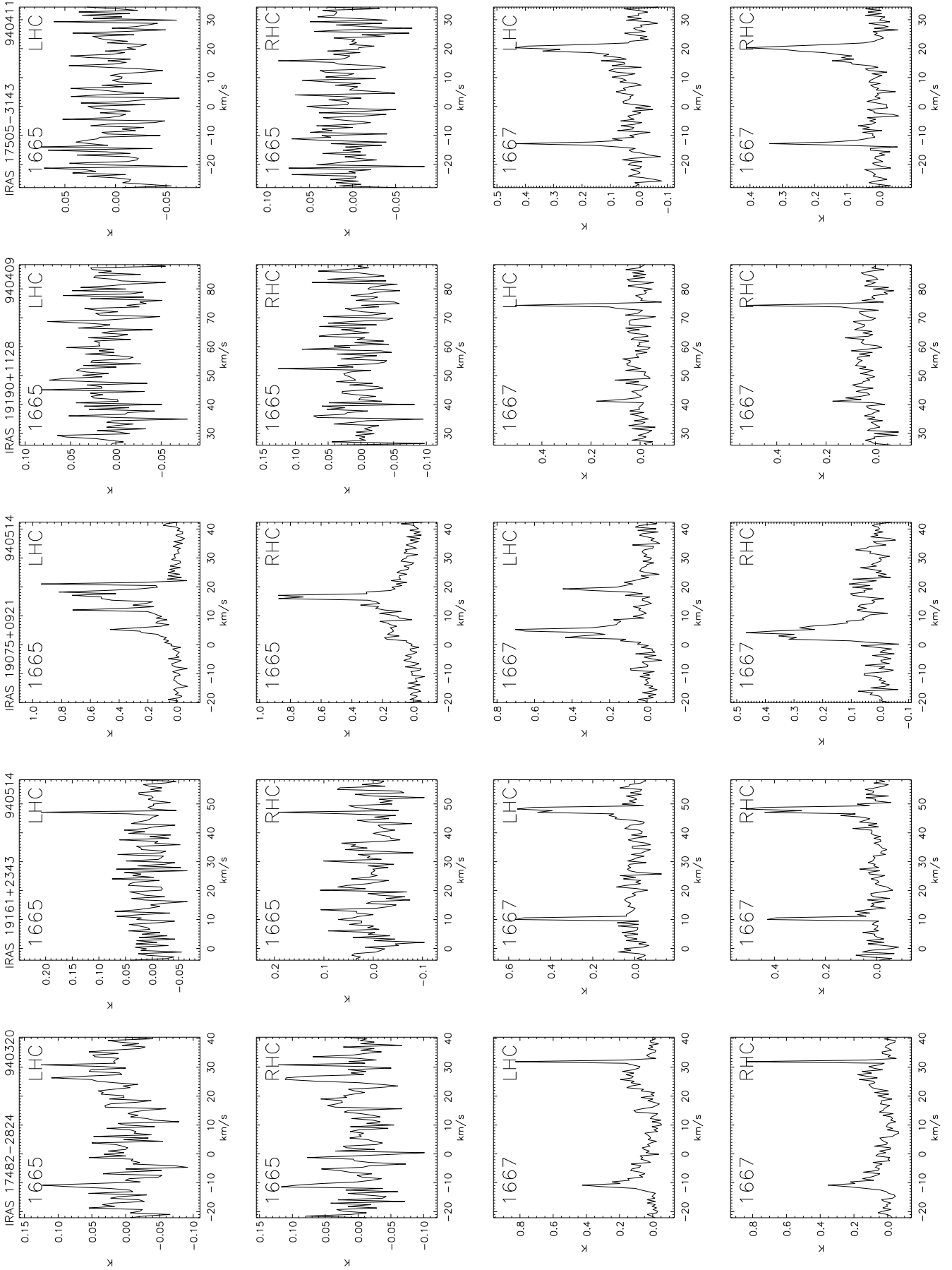
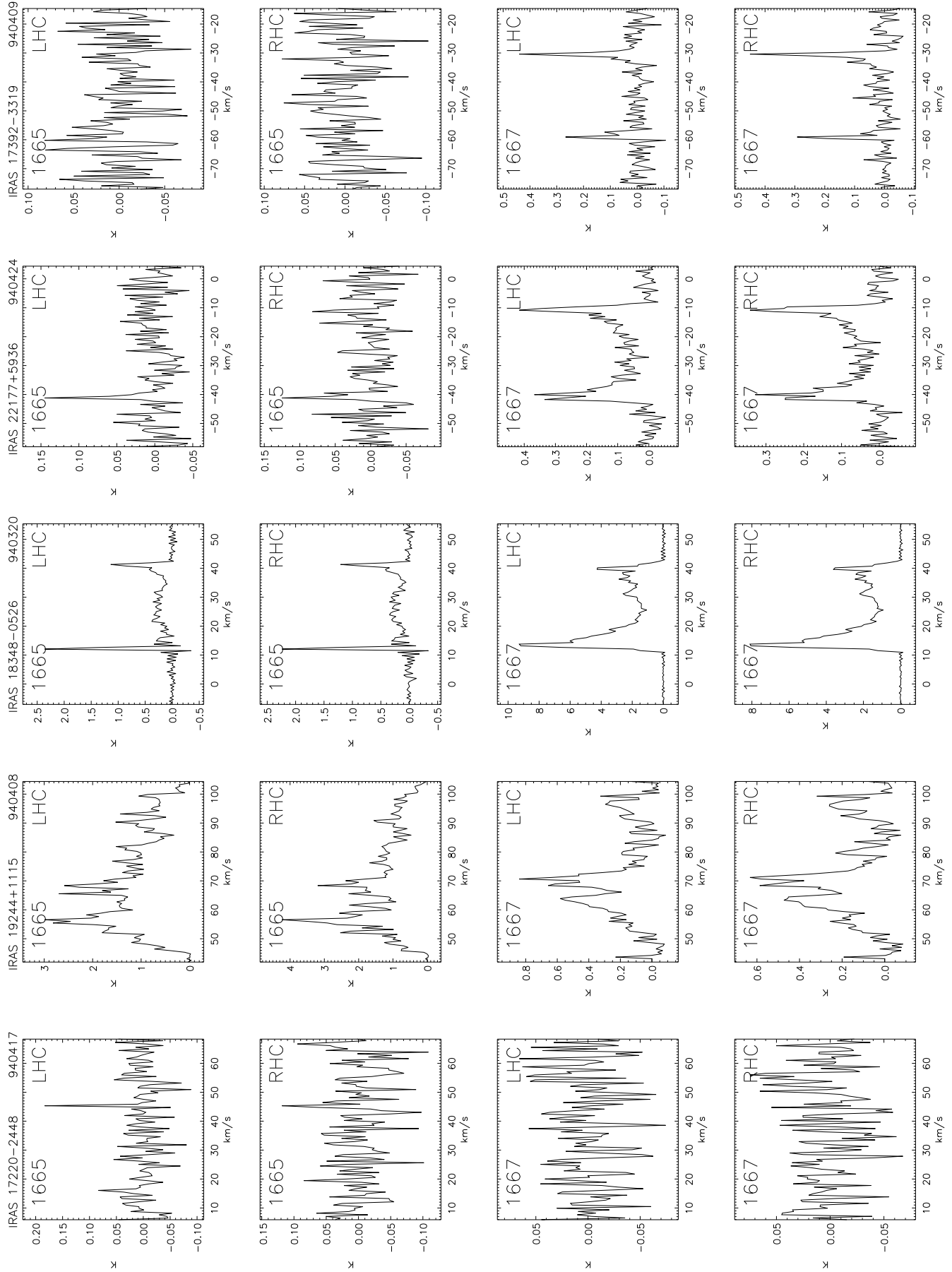
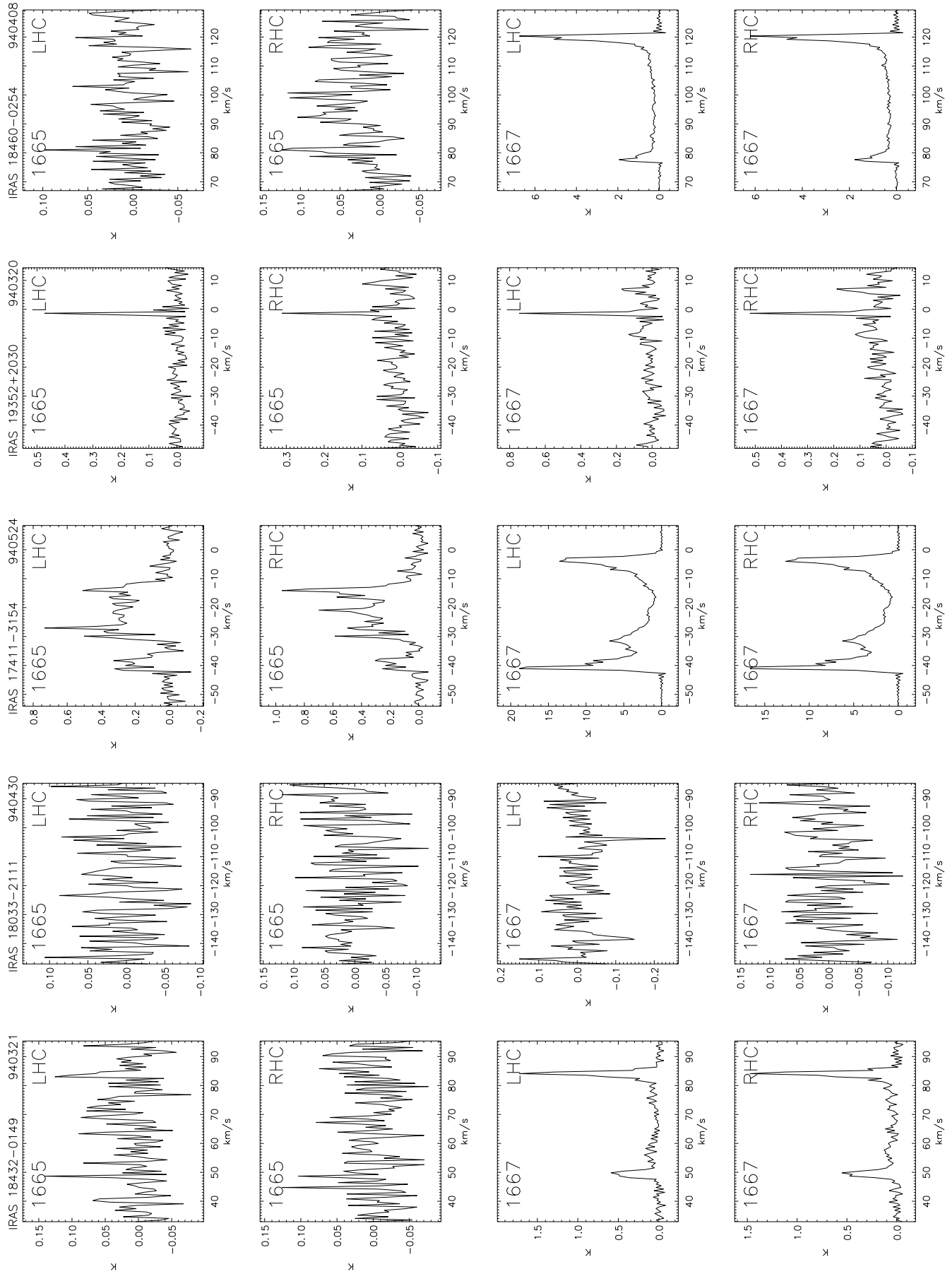
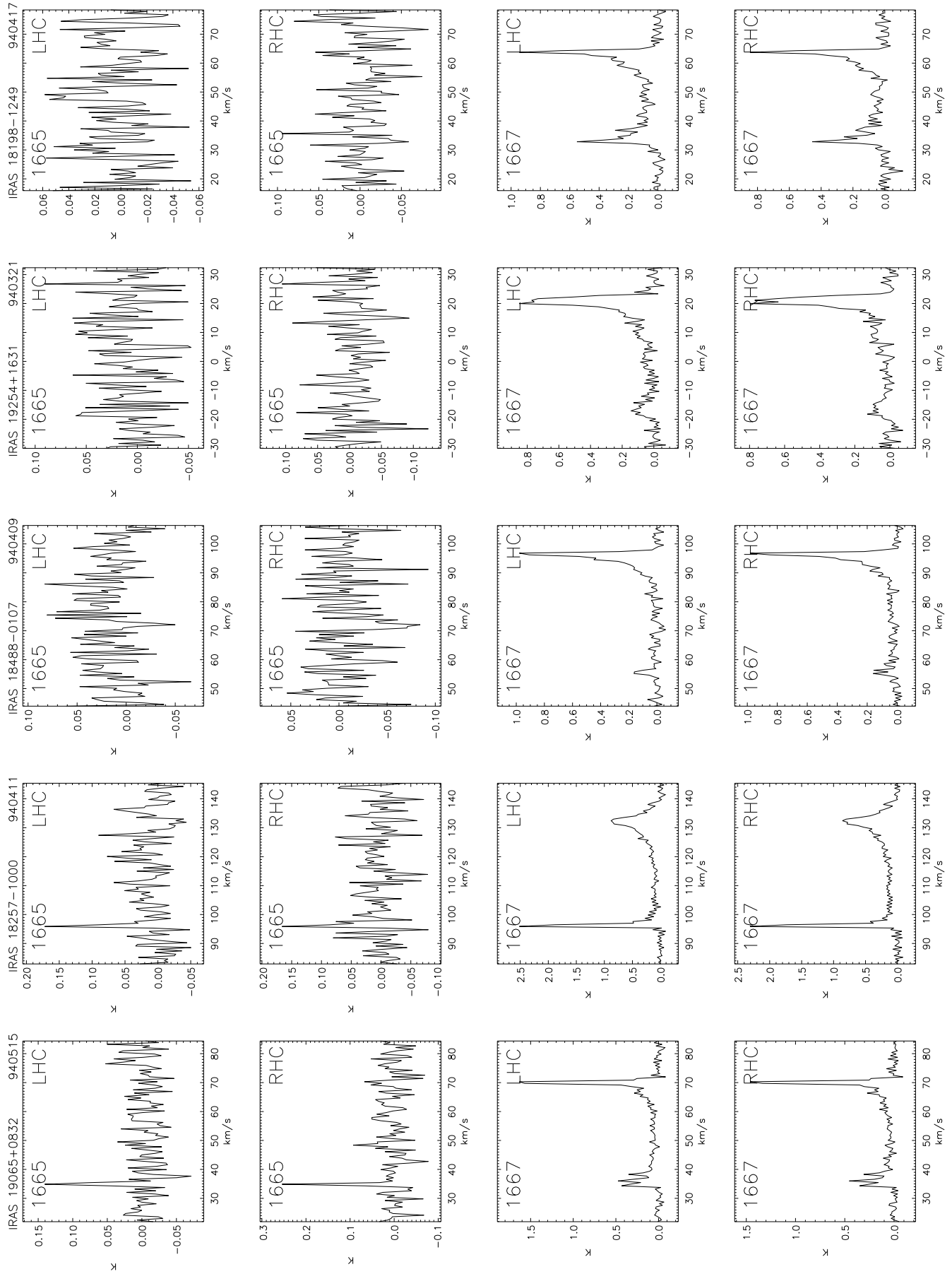


Fig. B.1. Spectra of the sources at 1665 and 1667 MHz in the left- (LHC) and right-handed (RHC) polarizations, appearing in increasing value of the $[25-12]$ colour index.

Fig. B.1. *cont.*

Fig. B.1. *cont.*

Fig. B.1. *cont.*

Fig. B.1. *cont.*

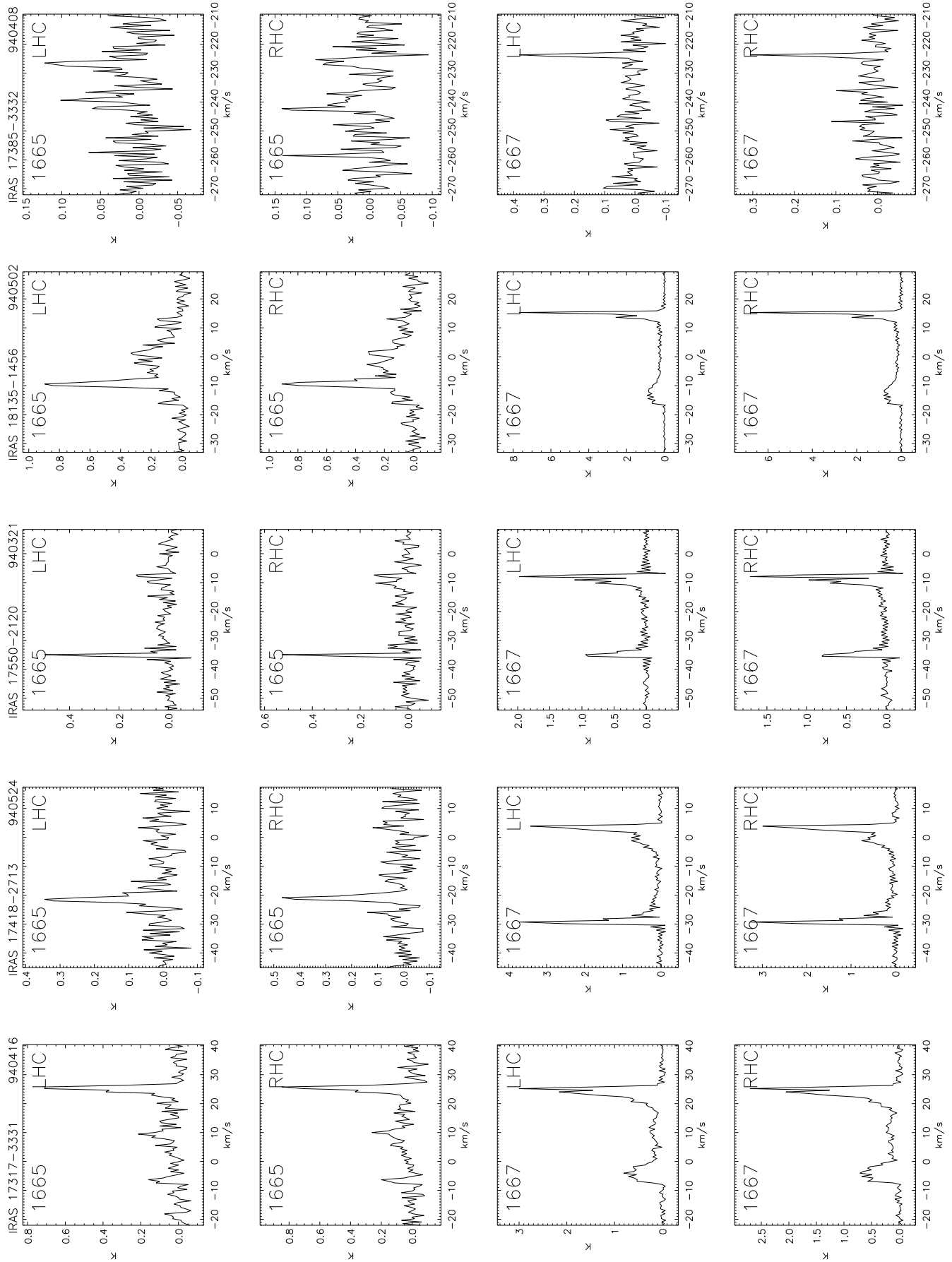
Fig. B.1. *cont.*

Table C.1. Results of the spectral measurements at 1665 MHz

Iras Name	ps ²	V _b ¹		V _r ¹		V _{int} ^{1,3}		V _* ¹	V _{exp} ¹	S _b ¹		S _r ¹		S _{int} ^{1,3}	
		L ¹	R ¹	L	R	L	R			L	R	L	R	L	R
15226–3603			
15262+0400		38.38	38.38	49.06	48.50			43.6	5.2	0.19	0.18	0.22	0.22		
20396–0826	c?p?	–51.13	–52.25			0.16	0.13
20547+0247		–75.75	–75.19					0.14	0.17		
22525+6033	p	–67.19?				–55.38	–55.38			0.19?				0.56	1.01
		–62.69		–41.87	–41.87	–53.12	–53.12	–52.3	10.4	0.41		0.25	0.34	0.66	0.38
		–61.00		–43.00		–49.75				0.39		0.43		0.36	
		–58.19	–58.19		–46.94	–48.62				0.26	1.52		0.29	0.29	
17482–2824		10.81	–11.38	30.81	30.81			9.9	20.9	0.12	0.12	0.13	0.12		
17505–3143			
19075+0921	p	5.25	2.44	21.00	17.06	12.00	8.06	11.4	7.6	0.46	0.19	0.94	0.87	0.72	0.22
		17.06	10.88			0.73	0.22
		18.19	13.69			0.82	0.34
19161+2343		47.13	47.13					0.21	0.19		
19190+1128			
19244+1115	p	56.56	56.56	99.31	98.13	65.56	58.81	77.6	21.1	2.99	4.23	1.05	0.97	2.69	2.56
		68.38	61.62			2.59	2.27
		76.81	65.56			1.60	2.29
		81.88	68.38			1.51	3.19
		90.31	76.25			1.53	1.68
		93.13	81.88			1.43	1.32
		90.88			1.56
		93.13			1.06
		96.50			0.95
17317–3331	ci	–6.31	–6.31	25.19	25.75	9.44	10.00	9.6	15.9	0.16	0.20	0.71	0.85	0.21	0.26
17220–2448		45.31	45.31					0.18	0.12		
17385–3332	ci	– ...	–258.25	–226.50	–225.37	–239.44	–242.15	–241.8	16.2	0.06?	0.14	0.12	0.09	0.10	0.14
17392–3319		–60.19?	...	–22.50?	...			–41.3?	18.8?	0.08?	...	0.07?	...		
17411–3154	p	–38.31	–38.31	–14.13	–14.13	–27.06	–20.88	–26.2	12.1	0.32	0.59	0.51	0.96	0.73	0.59
17418–2713	ci	–21.44	–20.88			0.34	0.47
17550–2120		–34.94	–34.94	–7.37	–7.37			–21.1	13.8	0.50	0.53	0.13	0.14		
18033–2111			
18135–1456	p	–16.19	–16.19	13.06	13.06	–9.44	–9.44	–1.6	14.6	0.16	0.15	0.16	0.18	0.90	0.91
		–2.13	–2.67			0.31	0.32
		1.25	1.81			0.33	0.31
		5.75	5.75			0.16	0.14
		10.25				0.18	
18198–1249			
18257–1000		95.87	95.87					0.17	0.17		
18348–0526	s	12.06	12.06	41.31	41.31	26.69	26.13	26.7	14.6	2.36	2.24	1.14	1.21	0.31	0.30
18432–0149		48.69	...	83.00	...			65.8	17.2	0.14	...	0.13	...		
18460–0254	s?	81.00	81.00		100.69?			0.10	0.12	0.12?	
18488–0107			
19065+0832		34.87	34.87					0.14	0.25		
19254+1631			
19352+2030	ci	8.75?	–1.37	–1.37			0.10?	0.47	0.31
22177+5936		–41.19	–41.19	...	–11.37?			–26.3?	14.9?	0.14	0.13	...	0.08?		

(1) **L**: left-handed polarization, **R**: right-handed polarization(2) **profile shape characteristic** :**p**: plateau profile, **c**: profile with only a single peak centred on the stellar velocity, **s**: “standard” profile with inter-peak signal, **ci**: profile with an internal component or group of components (well detached from the two standard peaks)(3) **V_{int}**: in one of the mentioned cases. When the profile is triple-peaked, the velocity of the component showing the strongest intensity (the latter might not be centred on the stellar velocity) is at least given. **S_{int}**: intensity corresponding to **V_{int}**

Table C.2. Results of the spectral measurements at 1667 MHz

Iras N	Name	ps ²	V _b ¹		V _r ¹		V _{int} ^{1,3}		V _* ¹	V _{exp} ¹	S _b ¹		S _r ¹		S _{int} ^{1,3}	
			km/s	R ¹	km/s	R	km/s	R			Jy	R	Jy	R	Jy	R
15226	–3603
15262	+0400	37.82	37.82	50.18	50.74	44.1	6.3	0.30	0.31	0.81	0.69
20396	–0826	...	–53.37?	–47.19	–47.19	50.3?	3.1?	0.07?	0.12	0.11
20547	+0247	–75.19	0.11
22525	+6033	–62.12	–57.63	–43.02	–46.39	–52.6	9.5	0.17	0.37	0.27	0.41
		–56.51	–53.70	–45.83	–47.51	0.48	0.75	0.16	0.37
		–53.70	...	–44.71	0.26	...	0.17
		–51.45	0.24
17482	–2824	ci	–10.79	–10.79	31.91	31.91	15.06	...	10.6	21.3	0.43	0.35	0.83	0.84	0.11	...
17505	–3143	...	–12.86	–12.86	20.29	20.29	3.7	16.6	0.38	0.33	0.48	0.46
19075	+0921	...	5.26	4.13	19.30	20.99	12.4	7.7	0.70	0.47	0.45	0.11
19161	+2343	...	10.02	10.02	48.23	48.23	29.1	19.1	0.57	0.43	0.56	0.51
19190	+1128	...	41.14	41.14	74.29	74.29	57.7	16.6	0.18	0.17	0.51	0.52
19244	+1115	ci	63.89	63.32	99.29	99.29	81.4	17.8	0.58	0.47	0.33	0.32
		70.63	71.19	79.62	0.84	0.63	0.23
17220	–2448
17317	–3331	s	–4.05	–4.05	25.17	25.17	10.00	10.00	–10.6	14.6	0.81	0.71	2.99	2.71	0.26	0.23
17385	–3332	ci?	–223.71	–223.71	–246.18?	–246.74?	0.38	0.31	0.09?	0.11?
		–236.07?	0.10?
17392	–3319	...	–59.05	–59.05	–30.39	–30.39	–44.7	14.3	0.27	0.29	0.44	0.45
17411	–3154	s	–41.10	–41.10	–4.02	–4.02	–22.66	–22.56	–22.6	18.5	18.87	16.67	13.55	12.67	1.93	1.69
17418	–2713	...	–29.29	–29.29	3.86	3.86	–12.7	16.6	3.72	3.28	3.43	2.99
17550	–2120	...	–34.92	–35.49	–7.95	–7.95	–21.6	13.6	0.94	0.81	1.98	1.71
18033	–2111
18135	–1456	s	–13.36	–13.36	15.29	15.29	1.24	1.24	1.0	14.3	0.91	0.79	7.67	6.82	0.30	0.15
18198	–1249	s	32.83	32.83	63.73	63.73	48.56	48.56	48.3	15.4	0.55	0.46	0.94	0.85	0.12	0.07
18257	–1000	s	95.90	95.90	132.42	132.42	114.44	114.44	114.2	18.3	2.52	2.31	0.88	0.86	0.11	0.13
18348	–0526	s	13.76	13.76	39.61	39.61	26.12	26.12	26.7	12.9	9.27	8.11	4.26	3.60	1.41	1.26
18432	–0149	...	49.83	49.83	84.10	84.10	67.0	17.1	0.59	0.55	1.72	1.47
18460	–0254	s	77.65	77.65	120.35	120.35	99.00	99.00	99.0	21.3	1.94	1.78	6.76	6.22	0.25	0.30
18488	–0107	...	55.21	55.21	96.79	96.79	76.0	20.8	0.16	0.16	0.98	0.98
19065	+0832	s?	36.02	36.02	70.29	70.29	53.44?	54.07?	53.2	17.1	0.47	0.46	1.63	1.46	0.09?	0.07?
19254	+1631	...	–17.10	–18.23	19.98	19.98	1.2	18.8	0.14	0.13	0.85	0.80
19352	+2030	ci	–8.70	–8.70	7.04	7.04	–1.39	–1.39	–0.8	7.9	0.13	0.12	0.17	0.19	0.75	0.52
22177	+5936	s	–40.05	–40.05	–10.83	–10.83	–26.00	26.56	–25.4	14.6	0.37	0.33	0.42	0.34	0.06	0.06

1,2 and 3 : same as for the previous table